

PROCEEDINGS OF SPIE

Current Developments in Lens Design and Optical Engineering XVII

**R. Barry Johnson
Virendra N. Mahajan
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Editors

**31 August – 1 September 2016
San Diego, California, United States**

Sponsored and Published by
SPIE

Volume 9947

Proceedings of SPIE 0277-786X, V. 9947

SPIE is an international society advancing an interdisciplinary approach to the science and application of light.

Current Developments in Lens Design and Optical Engineering XVII, edited by R. Barry Johnson,
Virendra N. Mahajan, Simon Thibault, Proc. of SPIE Vol. 9947, 994701 · © 2016 SPIE
CCC code: 0277-786X/16/\$18 · doi: 10.1117/12.2263009

Proc. of SPIE Vol. 9947 994701-1

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Author(s), "Title of Paper," in *Current Developments in Lens Design and Optical Engineering XVII*, edited by R. Barry Johnson, Virendra N. Mahajan, Simon Thibault, Proceedings of SPIE Vol. 9947 (SPIE, Bellingham, WA, 2016) Six-digit Article CID Number.

ISSN: 0277-786X
ISBN: 9781510602861 (electronic)
ISBN: 9781510602854

Published by

SPIE

P.O. Box 10, Bellingham, Washington 98227-0010 USA
Telephone +1 360 676 3290 (Pacific Time) · Fax +1 360 647 1445
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Contents

vii	<i>Authors</i>
ix	<i>Conference Committee</i>
xi	<i>Introduction</i>
xiii	<i>The advanced LIGO detectors in the era of first discoveries (Plenary Paper) [9960-200]</i>

SESSION 1 LENS DESIGN METHODOLOGY I

9947 02	A zoom lens from scratch: the case for number crunching (Invited Paper) [9947-1]
9947 03	Optical design and tolerancing of a hyperspectral imaging spectrometer [9947-2]
9947 04	Achrotech: achromat cost versus performance for conventional, diffractive, and GRIN components [9947-3]

SESSION 2 LENS DESIGN METHODOLOGY II

9947 06	A short history of nomograms and tables used for thermal radiation calculations (Invited Paper) [9947-5]
9947 07	Caustic surface produced by a plane wavefront refracted through positive bi-conic lenses [9947-6]
9947 08	Design investigation of a cost-effective dual-band (MWIR/LWIR) and a wide band optically athermalized application [9947-7]
9947 09	Novel microfluidic devices for Raman spectroscopy and optical trapping [9947-8]

SESSION 3 LENS DESIGN

9947 0A	Some methods for determining the limit of potential image quality of optical systems of various complexities using the database (Invited Paper) [9947-9]
9947 0B	Optical design of athermalised dual field of view zoom lens in long wave infrared (8μm - 12μm) spectral band using benefits of paraxial optics [9947-10]
9947 0C	Optical design of an ultrashort throw ratio projector with two freeform mirrors [9947-11]
9947 0E	Reduce volume of head-up display by image stitching [9947-13]

SESSION 4 OPTICAL TESTING AND ANALYSIS

- 9947 0F **Optical performance of a PDMS tunable lens with automatically controlled applied stress** [9947-14]
- 9947 0G **Fabrication of focus-tunable liquid crystal microlens array with spherical electrode** [9947-15]
- 9947 0H **Characterization of a tunable liquid-filled lens with minimum spherical aberration** [9947-16]
- 9947 0I **Low-cost automated system for phase-shifting and phase retrieval based on the tunability of a laser diode** [9947-17]

SESSION 5 APPLICATIONS AND ANALYSIS I

- 9947 0K **Stray light modeling of the James Webb Space Telescope (JWST) Integrated Science Instrument Module (ISIM)** [9947-19]
- 9947 0L **Design of a radiance meter with predicted size of source and distance effects** [9947-20]
- 9947 0M **Formation of the color image based on the vidicon TV camera** [9947-21]
- 9947 0N **Efficient 3M PBS enhancing miniature projection optics** [9947-22]
- 9947 0P **Studying the back-scattering of light for the development of acousto-optical filter with an improved spectral resolution** [9947-24]

SESSION 6 APPLICATIONS AND ANALYSIS II

- 9947 0Q **Accuracy and sensitivity analysis of the conical null-screen based corneal topographer** [9947-25]
- 9947 0T **Enhancement of light luminance film applied to the transparent display devices** [9947-28]
- 9947 0U **Retina projection using curved lens arrays** [9947-29]

POSTER SESSION

- 9947 0W **The thickness of DLC thin film affects the thermal conduction of HPLED lights** [9947-31]
- 9947 0X **Assembly aligning and measuring of a reflective telescope primary mirror** [9947-32]
- 9947 0Z **An overview of inverse solution expressions for Risley-prism-based scanner** [9947-34]
- 9947 10 **Evaluation of retinal illumination in coaxial fundus camera** [9947-35]

- 9947 14 **Corneal topography with conical null-screen for non-symmetric aspheric corneas** [9947-39]
- 9947 15 **Development of methods for accurate modeling of optical equipment for three-dimensional printing** [9947-40]
- 9947 19 **Fundamentals of concentric lens systems synthesis** [9947-44]
- 9947 1A **Composition variants for mirror high-aperture lens with compact design** [9947-45]
- 9947 1B **Focusing regions evolution under linear transformations in the boundary condition** [9947-46]
- 9947 1F **Three-dimensional ray tracing for refractive correction of human eye ametropies** [9947-50]
- 9947 1I **Infrared image acquisition system for vein pattern analysis** [9947-53]

Authors

Numbers in the index correspond to the last two digits of the six-digit citation identifier (CID) article numbering system used in Proceedings of SPIE. The first four digits reflect the volume number. Base 36 numbering is employed for the last two digits and indicates the order of articles within the volume. Numbers start with 00, 01, 02, 03, 04, 05, 06, 07, 08, 09, 0A, 0B...0Z, followed by 10-1Z, 20-2Z, etc.

Arellanes, Adan Omar, 0P
Armengol-Cruz, Victor de Emanuel, 14
Avendaño-Alejo, Maximino, 07
Beniot, Gilles J., 0N
Benítez, Pablo, 0C
Bezdidko, S., 0A
Burtseva, Anastasiia, 1A
Campos-García, Manuel, 0Q, 14
Castro Neto, Jarbas C., 10
Castro-Ortega, R., 11
Chan, Chia-Yen, 0X
Chaves, Julio, 0C
Chavushyan, Vahram, 0P
Chen, Yueh-Hao, 0T
Chien, Chao-Heng, 0T
Chien, Wei-Cheng, 0T
Chiu, Chuang-Hung, 0T
Chiu, Yi-Feng, 0E
Chuhlamov, Anton, 1A
Cossio-Guerrero, Cesar, 0Q, 14
Cruz-Felix, Angel S., 0F, 0H
de Coster, Diane, 09
De los Santos-García, S. I., 1B
de Matos, Luciana, 10
de Oliveira, André O., 10
Díaz-Gonzalez, G., 0H, 1F
Díaz-Uribe, José Rufino, 14
Dilworth, D. C., 02
Ding, Fujian, 08
Dirckx, Joris, 0I
Duerr, Fabian, 0C
Dunaeu, Vadim I., 0M
Ezhova, Kseniia, 19, 1A
Flügel-Paul, Thomas, 03
García-Guzman, A., 1B
Gross, Herbert, 03
Hernández Méndez, Arturo, 0F
Hsu, Ming Seng, 0W
Hsu, Ming-Ying, 0X
Huang, Jen Wei, 0W
Huang, Ting-Ming, 0X
Huang, Wei-Ming, 0G
Hurtado-Ramos, J. B., 0L
Irvin, Ryan G., 0K
Iturbe-Castillo, M. D., 1F
Iureva, Radda A., 0M
Jiménez-Barriga, N., 0H
Jiménez-Hernández, J. A., 1F
Johnson, R. Barry, 06
Juarez-Salazar, R., 0H, 1F
Kent, Susan, 0N
Kucukcelebi, Doruk, 0B
Le, John, 0N
Li, Anhu, 0Z
Lie, Chun-Chieh, 0X
Lin, Yu-Chuan, 0X
Liu, Chang, 03
Liu, Qing, 09
Maltseva, Nadezhda K., 0M
Martínez-Niconoff, G., 1B
McDowell, Erin, 0N
Miñano, Juan C., 0C
Mitiushkin, A. V., 15
Mohedano, Rubén, 0C
Moreno-Oliva, Iván, 07
Morgen, Daniel, 08
Morris, Jeffrey, 04
Mortenson, Dave, 0N
Munoz-Lopez, J., 0H, 1B
Nevitt, Timothy, 0N
Nie, Yunfeng, 0C
Osorio-Infante, Arturo I., 14
Ottevaere, Heidi, 09
Ouderkirk, Andrew, 0N
Padilla-Vivanco, A., 11
Ponce-Hernández, Osvaldo, 07
Reyes Pérez, Emilio R., 0F
Rivera-Ortega, Uriel, 0I
Rohrbach, Scott O., 0K
Román-Hernández, Edwin, 07
Saitgalina, A. K., 15
Santiago-Alvarado, Agustín, 0F, 0H, 1F
Seals, Lenward T., 0K
Shcherbakov, Alexandre S., 0P
Shyu, Feng Lin, 0W
Sigg, Daniel, xiii
Skelton, Dennis L., 0K
Solís-Villarreal, J., 11
Solorio-Leyva, J. C., 0L
Sparrold, Scott, 04
Stewart, Seán M., 06
Straif, Christoph, 03
Su, Guo-Dung J., 0E, 0G, 0U
Suárez-Romero, J. G., 0L
Sun, Wansong, 0Z
Tepichín-Rodríguez, Eduardo, 0F
Thienpont, Hugo, 09, 0C
Tochilina, Tatiana, 19

Tolstoba, N. D., 15
Torres, M. A., 1B
Toxqui-Quitl, C., 1I
Trujillo-Romero, F., 1F
Tymoshchuk, Irina, 19
Van Erps, Jürgen, 09
Vandendriessche, Stefaan, 04
Vervaeke, Michael, 09
Washer, Joe, 08
Willett, Stephen, 0N
Wolf, Greg, 04
Wong, Timothy, 0N
Wu, Kun-Huan, 0X
Yen, Hao-Ren, 0U
Yi, Wanli, 0Z
Yun, Zhisheng, 0N
Zeitner, Uwe D., 03
Zverev, Victor, 19, 1A

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- 1 Lens Design Methodology I
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- 5 Applications and Analysis I
Simon Thibault, Université Laval (Canada)
- 6 Applications and Analysis II
Ching-Cherng Sun, National Central University (Taiwan)

Introduction

We accepted 52 papers, included 24 posters, and organized them into six sessions. We had two sessions on Lens Design Methodology, one on Lens Design, another on Optical Testing and Analysis, and two on Applications and Analysis. Three of the sessions started with an invited paper each. Mexico contributed the most papers (19), followed by the Russian Federation (9), Taiwan (8), United States (6), and others. The Proceedings consist of 34 papers. There were some no-show papers. Manuscripts of a few were submitted, but they were not presented. These papers are not published in the Proceedings. All of the missing papers are from outside of the United States.

R. Barry Johnson
Virendra N. Mahajan
Simon Thibault

The Advanced LIGO Detectors in the Era of First Discoveries

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ABSTRACT

Following a major upgrade, the two advanced detectors of the Laser Interferometer Gravitational-wave Observatory (LIGO) held their first observation run between September 2015 and January 2016. The product of observable volume and measurement time exceeded that of all previous runs within the first 16 days of coincident observation. On September 14th, 2015 the Advanced LIGO detectors observed the transient gravitational-wave signal GW150914, determined to be the coalescence of two black holes, launching the era of gravitational-wave astronomy. We present the main features of the detectors that enabled this observation. At its core Advanced LIGO is a multi-kilometer long Michelson interferometer employing optical resonators to enhance its sensitivity. Four very pure and homogeneous fused silica optics with excellent figure quality serve as the test masses. The displacement produced by the event GW150914 was one 200th of a proton radius. It was observed with a combined signal-to-noise ratio of 24 in coincidence by the two detectors. At full sensitivity, the Advanced LIGO detectors are designed to deliver another factor of three improvement in the signal-to-noise ratio for binary black hole systems similar in masses to GW150914.

Keywords: LIGO, gravitational waves, black holes

1. INTRODUCTION

Einstein's theory of relativity predicts the existence of gravitational waves which are ripples in the space-time fabric and propagate at the speed of light.^{1,2} An indirect proof of their existence was first observed in the double pulsar system PSR B1913+16 by Hulse, Taylor and Weisberg.^{3,4} On September 14th, 2015, both Advanced LIGO detectors in the USA, H1 in Hanford, Washington and L1 in Livingston, Louisiana, made the first direct measurement of gravitational waves.⁵ The event, GW150914, was the merger of two black holes, with masses of $36_{-4}^{+5} M_{\odot}$ and $29_{-4}^{+4} M_{\odot}$, into a black hole of approximately $62_{-4}^{+4} M_{\odot}$ (90% confidence level).⁶ The equivalent of 3.0 solar masses of energy ($\simeq 5.4 \times 10^{47}$ J) was radiated in gravitational waves. The gravitational waves from this event, which occurred at a distance of $\simeq 410_{-180}^{+160}$ Mpc $\simeq 1.3 \times 10^9$ light years, changed the separation between the test masses by $\simeq 4 \times 10^{-18}$ m, or about one 200-th of a proton radius.

The Advanced LIGO detectors are 4 km long Michelson interferometers enhanced by multiple optical cavities.^{7,8} Each Michelson arm employs two test masses that form the arm cavities. Advanced LIGO came online in September 2015, after a major upgrade targeting a factor of 10 sensitivity improvement over initial detectors.^{9,10} While not yet at design sensitivity during their first observation run, they have already exceeded the strain sensitivity of the initial detectors across the entire frequency band. This has significantly increased the discovery potential and led to the discovery of the first black hole merger.¹¹⁻¹⁵

2. THE ADVANCED LIGO DETECTORS

The LIGO laboratory runs two observatories which house one Advance LIGO detector each. Figure 1 shows an areal view of the LIGO Hanford Observatory. The second observatory in Livingston is identical. Two or more observatories are required for coincident detection and to distinguish between signals of astrophysical origin and environmental disturbances.

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Figure 1. LIGO Hanford Observatory. The LIGO Hanford Observatory is located in eastern Washington. It houses one of the 4 km long Advanced LIGO detector.

The Advanced LIGO detectors are major upgrades to the initial detectors with much improved seismic isolation, quadruple pendulums, monolithic suspensions and electrostatic drives. The available laser power was also increased by a factor of 10, but for the first run only 20-25 W were used. Installation of the new detectors finished in June 2014 in Livingston and in January 2015 in Hanford. Figure 2 shows the progress in commissioning between end of installation and the begin of the first observational run. The displayed range represents the sky-averaged reach to detect the coalescence of a binary neutron with a signal-to-noise ratio of 8. The units are Mpc. These detectors, with a range of 70 Mpc, reach far beyond the local galaxy group and encompasses the entire Virgo supercluster. We also note that for heavier sources such as $30 M_{\odot}$ black hole mergers the range reaches beyond 1 Gpc.

Figure 3 shows a simplified optical layout of the Advanced LIGO detector. A 200 W pre-stabilized laser (left) is used as the main light source at a wavelength of 1064 nm (red beams).¹⁶ An electro-optic modulator (not shown) is used to impose phase modulated RF sidebands at 9 MHz (brown beams) and 45 MHz (blue beams). It is spatially cleaned by a suspended input mode cleaner (not shown),¹⁷ before being injected into the dual-recycled Michelson interferometer with arm cavities.¹⁸ Each arm cavity comprises an input test mass (ITMX and ITMY) and an end test mass (ETMX and ETMY). The beamsplitter (BS) separates the main beam into the X-arm and

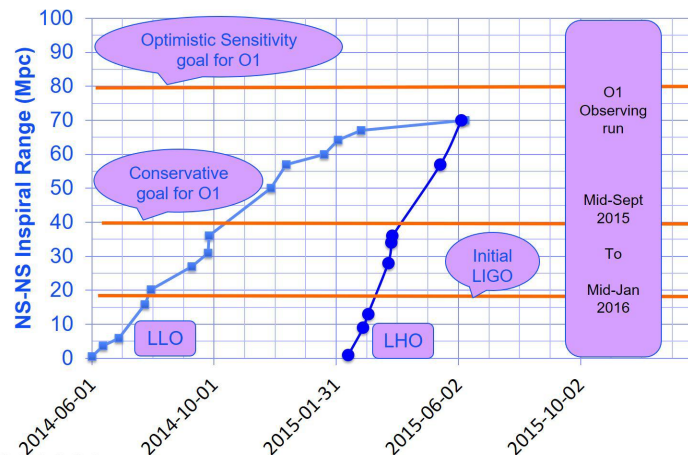


Figure 2. Commissioning progress. This graph shows the increase in binary neutron star reach as function of commissioning time, starting immediately after installation finished. The initial LIGO reach was surpassed within 2 months. The Hanford detector (LHO) finished installation about 8 months after the Livingston one (LLO), but caught up fast. Both detectors reached a sky-averaged range around 70 Mpc for binary neutron star inspirals with a single detector SNR of 8.

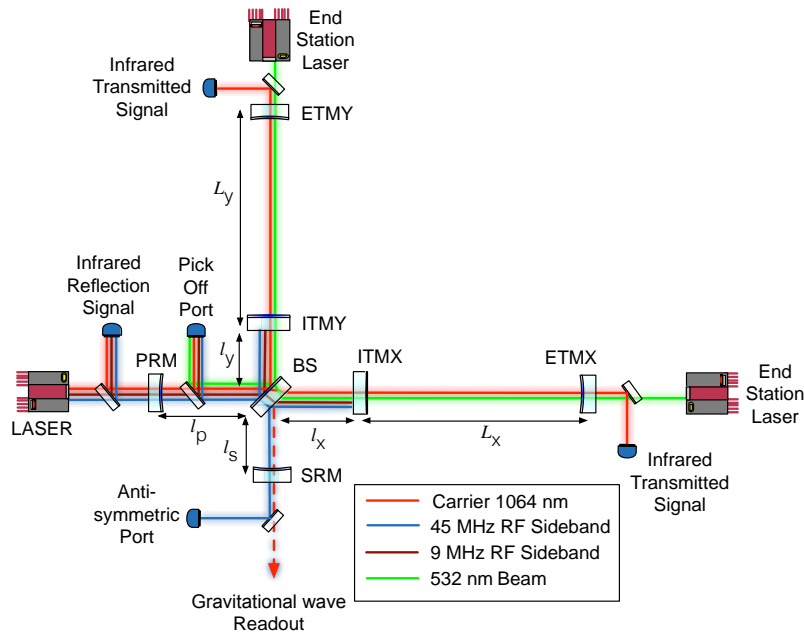


Figure 3. Simplified optical layout. See text for details.

the Y-arm. Two coupled cavities are formed around the Michelson vertex by the power recycling mirror (PRM) and by the signal recycling mirror (SRM), comprising the dual recycling configuration.^{19–22} Photodetectors are mounted in transmission of the arm cavities, at the antisymmetric port, in reflection of the interferometer and at the power recycling pick-off. Each of the multiple coupled optical cavities need to be locked to an exact multiple of the laser wavelength for operations.²³ Furthermore, the beamsplitter is placed at a dark fringe. An RF scheme is used to sense most of the length and angular degrees-of-freedom.^{24,25} An output mode cleaner (not shown) is mounted at the anti-symmetric port to suppress unwanted light frequencies. The photodetectors in transmission of the output mode cleaner are used for the gravitational wave readout. The gravitational wave readout is sensed using a DC scheme by shifting the arm cavities slightly off resonance.^{26,27} Doubled Nd:YAG lasers running at a wavelength of 532 nm are installed behind the end test masses (green beams) and are used to lock the arm cavities independent of the main laser source,²³ but only during the lock acquisition process and not during observations.

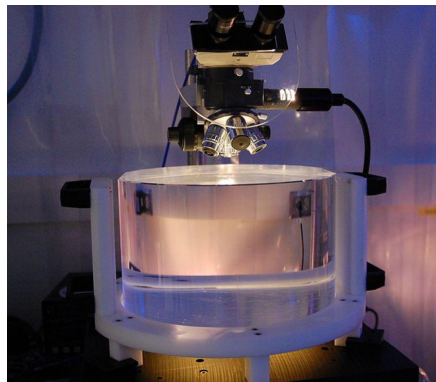


Figure 4. Test mass.

High power operation poses special challenges with thermal lensing, angular instabilities driven by radiation pressure²⁸ and parametric instabilities.²⁹ The system is typically locked at a lower power of 2 W, where these effects are negligible, and then powered up. Thermal lenses are corrected with a thermal compensation system³⁰ which comprises two CO2 lasers for central heating and four ring heaters for annular heating of the test masses. Parametric instabilities are driven by radiation pressure and the interaction between the test mass acoustic modes and the higher order optical modes. They are suppressed using active controls which applies damping filters. This system needs to engage during the power up to prevent run-away oscillations which break the lock. The higher the optical power in the interferometer the larger number of parametric instabilities are present. This makes the problem relatively easy at 20 W but hard at 200 W. Angular instabilities due to radiation pressure also depend on the power in the interferometer and become harder to damp at higher power. The current strategy requires angular servos with a higher bandwidth than the frequencies of the angular instabilities. However, increasing the bandwidth of the angular servos introduces additional noise in the gravitational wave readout, unless there is a sharp cut-off filter in the response.

A picture of a test mass in the metrology lab is shown in Figure 4. The Advanced LIGO test masses are very pure and homogeneous fused silica mirrors of 34 cm diameter, 20 cm thickness and 40 kg mass. It is critical that the test masses be free from sources of displacement noise, such as environmental disturbances from seismic noise,³¹ or thermally driven motion. To reduce the effects of ground vibrations, the test masses are suspended by multi-stage pendulums,³² thus acting as free masses well above the lowest pendulum resonance frequency of 0.4 Hz. Monolithic fused silica fibers³³ are incorporated at the bottom stage to decrease suspension thermal noise,³⁴ which limits the sensitive band. The Advanced LIGO test masses require about 10 orders of magnitude suppression of ground motion above 10 Hz. The multi-stage pendulum system attenuates the ground motion by seven orders of magnitude. It is mounted on an actively controlled seismic isolation platform which provides three orders of magnitude of isolation of its own.^{35,36} Moreover, these platforms are used to reduce the very large displacements produced by tidal motion and microseismic activity. Tidal forces can produce displacements up to several 100 μm over a multi-kilometer baseline on time scales of hours. The dominant microseismic activity is driven by ocean waves. The resulting ground motion can be as large as several μm at frequencies around 0.15 Hz—even far inland.

The entire test mass assembly including the suspension system and part of the seismic isolation system resides inside an ultra-high vacuum system, with pressures typically below 1 μPa over the 10,000 m^3 volume, to prevent acoustic shorting of the seismic isolation systems and to reduce Rayleigh scattering in the optical readout.

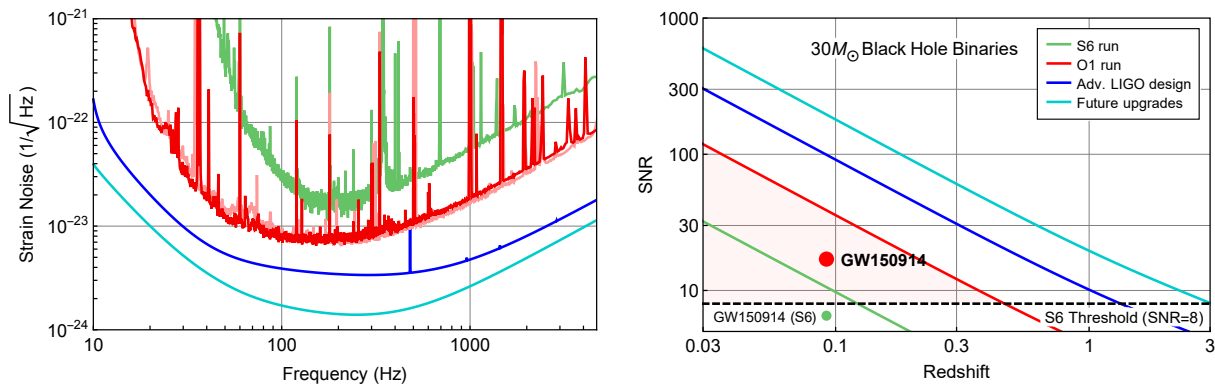


Figure 5. Sensitivity during the first observation run.⁷ The left plot shows the strain sensitivity during the first observation run (O1) of the Advanced LIGO detectors and during the last science run (S6) of the initial LIGO detectors. The O1 strain noise curve is shown for H1 (dark red) and L1 (light red); the two detectors have similar performance. The Advanced LIGO design sensitivity as well as a possible future upgrade³⁷ are shown to highlight the discovery potential in the coming years. The right plot shows the single detector signal-to-noise ratio (SNR) under optimal orientation as function of redshift z —for two merging black holes with mass $30 M_{\odot}$ each. GW150914 was not optimally orientated and was detected with a single detector SNR of 13 to 20 at $z = 0.09$; this event would not have been detected in initial LIGO.

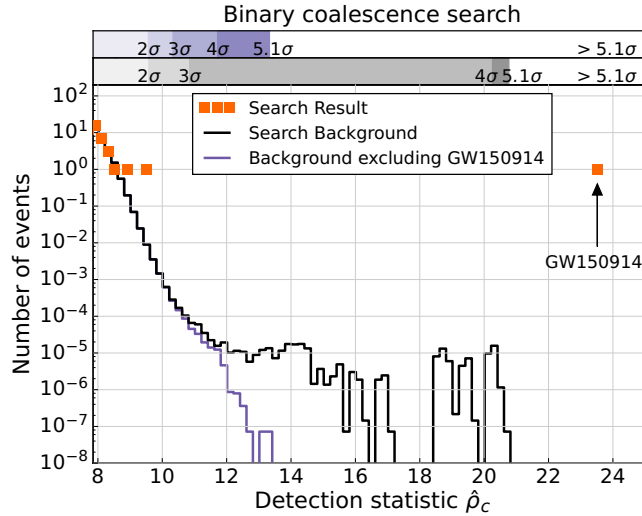


Figure 6. First detection.⁵ The foreground events are denoted by the red squares. The background is estimated using time shifted data from the two detectors. GW150914 is the most significant event in the first 16 days of coincident observation. When the event is removed from the data (blue line), the remaining background is clean. If we look at the search background including GW150914, we see a tail towards higher detection statistics which consist of the event being detected in one of the detectors in coincidence with a slightly elevated background in the other detector.

3. THE FIRST DETECTION

The sensitivity of the Advanced LIGO detectors during the first observation run³⁸ is shown in Figure 5. The sensitivity is significantly better than what was achieved during the last science run of the initial detectors. This is especially true for frequencies below 200 Hz which are most important for black hole mergers. It is important to remember that the observed volume and thus the number of possible sources scales with the third power of the sensitivity. Reaching the Advanced LIGO design sensitivity will be crucial in making black hole mergers a routine occurrence in astrophysical observations.

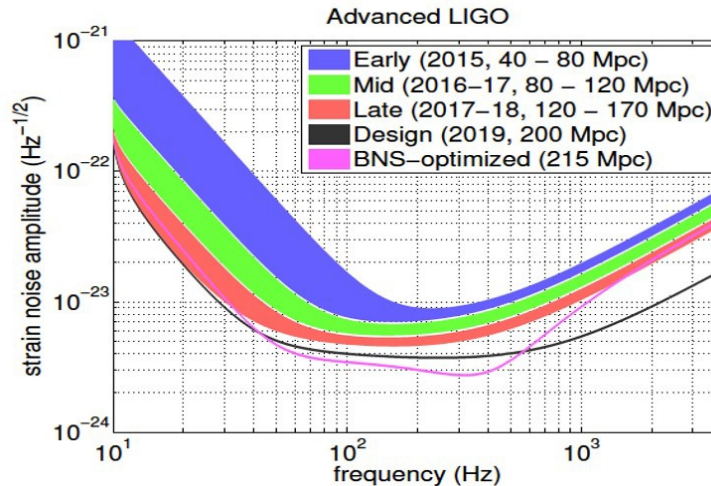


Figure 7. Planned observation periods.³⁹ The first observation run is depicted by the blue band. The second observation run is depicted by the green band and will take place in the second half of 2016 and early 2017. A third observation run is currently being planned for the late 2017 and 2018 time frame.

The first black hole merger event GW150914 was observed with a very high confidence level, see Figure 6. The event was detected by both LIGO observatories in coincident. It also constitutes the single largest such event in either of the two detectors. We estimate the background rate by analysing the time shifted data streams between the two observatories.

4. FUTURE PLANS

Future plans are depicted in Figure 7. With the successful completion of the first observation run, the commissioning for the second observation run is in progress. The second observation run is currently being planned to start in September 2016.

The second observation run, LIGO will be joined by the Virgo detector⁴⁰ which is located near Pisa, Italy. Having a third detector in the network is important for sky localization. With the LIGO detectors alone, the sky localization 90% confidence region can encompass several 100 square degrees. With the addition of the Virgo detector this potentially can be reduced to several square degrees.

The currently operating detectors are LIGO Hanford and LIGO Livingston as well as the GEO600⁴¹ detector. However, GEO600 is not sensitive enough to detect black hole mergers. The Virgo detector and the Japanese KAGRA detector⁴² are under construction and will come online over the next year or two. The third Advanced LIGO detector is planned to be installed in India.⁴³ Once the global network is active, we will be able to achieve good localization on any point in the sky as well as provide good coverage.

5. CONCLUSIONS

The LIGO Scientific Collaboration and the Virgo Collaboration have made the first detection of a black hole merger. This constitutes the first direct detection of gravitational waves on Earth. This is the beginning of the era of gravitational wave astrophysics with the promise of many more detections in the near future.

ACKNOWLEDGMENTS

The authors gratefully acknowledge the support of the United States National Science Foundation (NSF) for the construction and operation of the LIGO Laboratory and Advanced LIGO as well as the Science and Technology Facilities Council (STFC) of the United Kingdom, the Max-Planck-Society (MPS), and the State of Niedersachsen/Germany for support of the construction of Advanced LIGO and construction and operation of the GEO600 detector. Additional support for Advanced LIGO was provided by the Australian Research Council. The authors gratefully acknowledge the Italian Istituto Nazionale di Fisica Nucleare (INFN), the French Centre National de

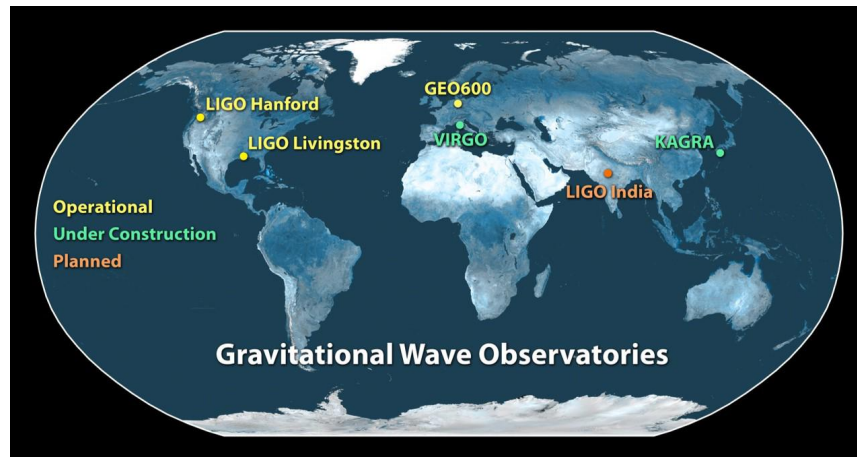


Figure 8. Detector network. This world map depicts the location of the planned network of ground based gravitational wave antennas.

la Recherche Scientifique (CNRS) and the Foundation for Fundamental Research on Matter supported by the Netherlands Organisation for Scientific Research, for the construction and operation of the Virgo detector and the creation and support of the EGO consortium. The authors also gratefully acknowledge research support from these agencies as well as by the Council of Scientific and Industrial Research of India, Department of Science and Technology, India, Science & Engineering Research Board (SERB), India, Ministry of Human Resource Development, India, the Spanish Ministerio de Economía y Competitividad, the Conselleria d’Economia i Competitivitat and Conselleria d’Educació, Cultura i Universitats of the Govern de les Illes Balears, the National Science Centre of Poland, the European Commission, the Royal Society, the Scottish Funding Council, the Scottish Universities Physics Alliance, the Hungarian Scientific Research Fund (OTKA), the Lyon Institute of Origins (LIO), the National Research Foundation of Korea, Industry Canada and the Province of Ontario through the Ministry of Economic Development and Innovation, the Natural Science and Engineering Research Council Canada, Canadian Institute for Advanced Research, the Brazilian Ministry of Science, Technology, and Innovation, Fundação de Amparo à Pesquisa do Estado de São Paulo (FAPESP), Russian Foundation for Basic Research, the Leverhulme Trust, the Research Corporation, Ministry of Science and Technology (MOST), Taiwan and the Kavli Foundation. The authors gratefully acknowledge the support of the NSF, STFC, MPS, INFN, CNRS and the State of Niedersachsen/Germany for provision of computational resources.

This document has been assigned the LIGO Laboratory document number LIGO-P1600182.

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